

A polarization survey of bright extragalactic AT20G sources

M. Massardi,¹* S. G. Burke-Spolaor,² T. Murphy,^{3,4} R. Ricci,¹ M. López-Caniego,⁵
M. Negrello,⁶ R. Chhetri,^{2,7} G. De Zotti,^{6,8} R. D. Ekers,² R. B. Partridge⁹
and E. M. Sadler⁵

¹INAF–Istituto di Radioastronomia, Via Gobetti 101, I-40129 Bologna, Italy

²Australia Telescope National Facility, CSIRO, PO Box 76, Epping, NSW 1710, Australia

³School of Physics, University of Sydney, Sydney, NSW 2006, Australia

⁴School of Information Technologies, University of Sydney, NSW 2006, Australia

⁵Instituto de Física de Cantabria (CSIC-UC), Avda. los Castros s/n, E-39005 Santander, Spain

⁶INAF–Osservatorio Astronomico di Padova, vicolo dell’ Osservatorio 5, I-35122 Padova, Italy

⁷Department of Astrophysics and Optics, School of Physics, University of New South Wales, NSW 2052, Australia

⁸SISSA, Via Bonomea 265, I-34136 Trieste, Italy

⁹Haverford College Astronomy Department, 370 Lancaster Avenue, Haverford, PA 19041 USA

Accepted 2013 September 10. Received 2013 August 28; in original form 2013 July 26

ABSTRACT

We present polarization data for 180 extragalactic sources extracted from the Australia Telescope 20 GHz (AT20G) survey catalogue and observed with the Australia Telescope Compact Array during a dedicated, high sensitivity run ($\sigma_P \sim 1$ mJy). For the sake of completeness, we extracted the polarization information for seven extended sources from the 9 yr *Wilkinson Microwave Anisotropy Probe* co-added maps at 23 GHz. The full sample of 187 sources constitutes a $\simeq 99$ per cent complete sample of extragalactic sources brighter than $S_{20\text{GHz}} = 500$ mJy at the selection epoch with declination $\delta < -30^\circ$. The sample has a 91.4 per cent detection rate in polarization at ~ 20 GHz (94 per cent if considering the subsample of point-like sources). We have measurements also at 4.8 and 8.6 GHz within ~ 1 month of the 20 GHz observations for 172 sources to reconstruct the spectral properties of the sample in total intensity and in polarization: 143 of them have a polarization detection at all three frequencies.

We find that there is no statistically significant evidence of a relationship either between the fraction of polarization and frequency or between the fraction of polarization and the total intensity flux density. This indicates that Faraday depolarization is not very important above 4.8 GHz and that the magnetic field is not substantially more ordered in the regions dominating the emission at higher frequencies (up to 20 GHz). We estimate the distribution of the polarization fraction and the polarized flux density source counts at ~ 20 GHz.

Key words: polarization – techniques: polarimetric – surveys – galaxies: active – radio continuum: galaxies.

1 INTRODUCTION

The study of the properties of radio source populations above 10 GHz has progressed greatly in recent years, fostered by cosmic microwave background (CMB) observation campaigns that require an accurate understanding of the contamination of the CMB signal by foreground sources. Extragalactic radio sources are the dominant contaminant on angular scales smaller than 30 arcmin, both in total intensity and in polarization at frequencies of up to $\simeq 100$ –200 GHz (Toffolatti et al. 1998, 1999; de Zotti et al. 1999). An accurate determination of their emission is therefore important to get clean CMB

maps and is absolutely critical for the detection of the extremely weak cosmological B-mode polarization (see Tucci & Toffolatti 2012, for a recent review).

Extending the characterization of the polarization properties of radio sources to high frequencies is interesting per sé, as it provides information about the physics of the emission process. In compact, Doppler boosted sources that dominate the high-frequency population at bright flux density levels, the emission at higher and higher frequencies mostly arises from synchrotron, self-absorbed, knot-like structures in the relativistic jet closer and closer to the active nucleus (e.g. Blandford & Königl 1979). It has been argued that the ordering of magnetic fields should increase in the inner regions, and as a consequence, the polarization degree increases (Tucci et al. 2004).

*E-mail: massardi@ira.inaf.it

However, the polarization properties of high-frequency extragalactic contaminants are still poorly constrained by observations. Most current estimates rely on extrapolations from low-frequency samples; the NVSS at 1.4 GHz (Condon et al. 1998) still constitutes the largest sample of sources surveyed both in total intensity and polarization. Extrapolations are affected by large uncertainties since a complex combination of effects must be considered. This includes intrabeam effects and bandwidth depolarization, in addition to intrinsic frequency-dependent changes. The propagation of the radiation through diffuse plasma screens between the source and the observer can cause depolarization and rotation of the polarization angle. These effects are difficult to isolate observationally, although we can benefit from the inverse square frequency dependence of the latter effect.

Because the polarized signal in extragalactic objects is typically a few per cent of the total intensity signal, deep surveys are necessary to collect statistically significant samples. But high frequency, deep surveys are time consuming for diffraction-limited, ground-based telescopes. This has motivated sensitive high-frequency polarization measurements of source samples usually selected from surveys at $\lesssim 5$ GHz (Klein et al. 2003; Ricci et al. 2004; Jackson et al. 2007; Agudo et al. 2010; see Tucci et al. 2004, for a summary of earlier polarization surveys).

The study of polarization properties of complete samples selected at ≥ 20 GHz has become possible thanks to the *Wilkinson Microwave Anisotropy Probe* (*WMAP*) all-sky survey with a completeness limit of $\simeq 1$ Jy at 23 GHz (Argüeso, González-Nuevo & Toffolatti 2003; de Zotti et al. 2005; Wright et al. 2009; Gold et al. 2011) and to the Australia Telescope 20 GHz (AT20G) survey (Murphy et al. 2010; Hancock et al. 2011) that has covered the full southern sky with 91 per cent completeness above 100 mJy and 79 per cent completeness above 50 mJy in regions south of declination -15° (north of -15° the catalogue completeness is lower between 14 and 20 h in right ascension). Table 1 lists some details on the AT20G and *WMAP* related samples that include data in polarization. Larger samples of radio sources selected at higher frequencies are being provided by the *Planck* mission (Planck Collaboration 2011a,b, 2013).

Multiteradian samples of high-frequency selected polarized sources are also important for identifying suitable calibrators for CMB polarimetric experiments and upcoming millimetre-wave telescopes. Pictor A has been identified as a suitable extragalactic polarization calibrator for the *Planck* Low Frequency Instrument (LFI) because of its position (in the region of the ecliptic pole where the satellite scans \sim once per minute) and its lack of variability from the hotspots that dominate the polarized signal (Burke-Spolaor et al. 2009).

The polarization of *WMAP* sources has been investigated by López-Caniego et al. (2009) using *WMAP* data; 14 extragalactic sources were significantly detected in polarization. Follow-up observations of a complete sample of 203 *WMAP* sources were carried

out with the Very Large Array (VLA) by Jackson et al. (2010); polarized emission was detected for 123, 169 and 167 at 8.4, 22 and 43 GHz, respectively.

Sadler et al. (2006) presented polarization measurements for a sample of 173 AT20G sources brighter than $S_{20\text{GHz}} = 100$ mJy; 129 ($\simeq 75$ per cent) were detected at 20 GHz, with a median fractional polarization of 2.3 per cent. Massardi et al. (2008) discussed the polarization properties of the AT20G bright sample ($S_{20\text{GHz}} \geq 500$ mJy), finding 213 $\geq 3\sigma$ polarization detections at 20 GHz out of a total of 320 sources ($\simeq 67$ per cent), with a median fractional polarization of 2.5 per cent at 20 GHz. The spectral indices in total intensity and in polarization were found to be similar on average, but there were several sources for which the spectral shape of the polarized emission is substantially different from the spectral shape in total intensity. The full AT20G catalogue (Murphy et al. 2010; Massardi et al. 2011a) includes the 20 GHz polarized intensity for 768 sources, 467 of which also have simultaneous polarization detections at 5 and/or 8 GHz, out of a total of 5890 sources. The detection limit is defined as $\max(3\sigma, 0.01S_{20\text{GHz}}, 6\text{ mJy})$.

Sajina et al. (2011) obtained polarization measurements with the VLA at 4.86, 8.46, 22.46 and 43.34 GHz of 159 out of the $\simeq 200$ AT20G radio galaxies with $S_{20\text{GHz}} \geq 40$ mJy in an equatorial field of the Atacama Cosmology Telescope survey; polarized flux was detected at >95 per cent confidence level (c.l.) for 141, 146, 89 and 59 sources, from low to high frequencies. The measured polarization fractions are typically <5 per cent, although in some cases they are measured to be up to $\simeq 20$ per cent. For sources with detected polarized flux in all four bands, about 40 per cent of the sample, the polarization fractions typically increase with frequency. This trend is stronger for steeper spectrum sources as well as for the lower flux density sources.

The conclusions of all the polarization studies of complete samples selected at high frequencies are limited by the moderate detection rates. To overcome this limitation we have performed dedicated high-sensitivity polarization observations of a complete AT20G bright source subsample achieving a $\simeq 93$ per cent detection rate. In addition to allowing a more thorough investigation of the polarization properties of the high-frequency radio source populations, this sample constitutes a legacy data set for polarization studies in the Southern hemisphere. In particular, this sample could help in the definition of calibrator source lists for facilities working in the millimetric bands, like the Atacama Large Millimeter Array (ALMA) that, in its final configuration, will observe from 30 to 950 GHz in both total intensity and polarization.

This paper is structured as follows. In Section 2, we describe the sample selection and the observing strategy, and in Section 3 the data reduction techniques. For the sake of completeness, we included the polarized flux density measurements in this analysis as well as limits for the extended sources as extracted from low-resolution 9 yr *WMAP* maps described in Section 4. The final sample

Table 1. AT20G and *WMAP* related catalogues that include data in polarization.

References	Frequency(GHz)	No. of sources	Notes
Massardi et al. (2008) AT20G-BSS	4.8, 8.6, 20	320	AT20G bright sample
López-Caniego et al. (2009)	23, 33, 41	22	Polarization detection in <i>WMAP</i> map
Jackson et al. (2010), Battye et al. (2011)	8.4, 22, 43	230	<i>WMAP</i> sources
Sajina et al. (2011)	4.8, 8.4, 22, 43	159	Equatorial AT20G sources
Murphy et al. (2010), Massardi et al. (2011a) AT20G	4.8, 8.6, 20	5890	AT20G 91 per cent complete sample above 100 mJy
Burke-Spolaor et al. (2009)	18	9	Extended sources in the Southern hemisphere
Current paper	4.8, 8.6, 18	193	Complete sample above 500 mJy

thus includes 187 sources. In Section 5, we present the analysis of their spectral and polarization properties. In Section 6, we derive the source counts as a function of the polarized flux density. In Section 7, we present a selection of candidate polarization calibrators for millimetric wavelength studies. The main conclusions are summarized in Section 8.

2 SOURCE SELECTION AND OBSERVATIONS

The selection of the sample was based on the list of confirmed AT20G sources available at the epoch of our observations (2006 October). We selected all objects with flux density $S_{20\text{GHz}} > 500$ mJy and declination $\delta < -30^\circ$, excluding the Galactic plane region ($|b| \leq 1.5^\circ$) and the Large Magellanic Cloud (LMC) region (inside a circle of 5.5° radius centred at $\alpha = 05:23:34.7$ and $\delta = -69:45:22$). This resulted in a complete sample of 189 sources.

Nine of them were found, with the aid of low-frequency radio imaging surveys (PMN, Griffith et al. 1994, 1995; SUMSS, Mauch et al. 2003), to be very extended. These were observed with the mosaic mode at 20 GHz by Burke-Spolaor et al. (2009). Flux densities integrated over the whole source are available for five of them. For the remaining four objects the measured integrated flux densities, if available, refer to subregions. For these four objects we decided to extract the integrated flux densities from the low-resolution *WMAP* maps as described in Section 4.

It should be noted that the flux densities reported in the final published AT20G catalogue may be slightly different from those given in the preliminary 2006 version used for our source selection. The reason is that in the case of sources observed more than once (which is most likely the case for bright sources such as those considered here) the highest quality observation was listed in the final AT20G catalogue, as discussed by Murphy et al. (2010). Because of variability this has the effect of moving some sources above or below the adopted flux density threshold ($S_{20\text{GHz}} = 500$ mJy). Since there are more sources below than above the threshold, there are more sources moving up than moving down, and we end up with slightly more sources above threshold in the final catalogue than in the version we have used. Hence, there are 214 sources listed in the final AT20G catalogue with declination $\delta < -30^\circ$ and flux density above our chosen threshold.

On the other hand, of the 180 sources with good quality flux density in our Australia Telescope Compact Array (ATCA) observations (as discussed below) only 165 still have $S_{20\text{GHz}} \geq 500$ mJy in the final catalogue and only 145 were found to be above this threshold in our 2006 October observations. However, despite variability, the final sample, which is reasonably complete at the selection epoch, is representative of the bright 20 GHz population as a whole and can be used to assess statistical properties of this population.

2.1 Observations in the 20 GHz band

Observations were taken on 2006 October 1 using the most compact hybrid configuration of ATCA, ‘H75’, excluding the data from the farthest antenna. The longest baseline of this configuration is 75 m, and its T-shape ensures adequate Fourier coverage for snapshots taken on a relatively small range of hour angles and at high elevation.

Although the linear feeds of the ATCA somewhat complicate the polarization calibration procedure, the array has several inherent advantages for polarization experiments. The on-axis receivers of the telescope introduce relatively low amounts of instrumental polarization, while all antennas are fitted with a noise diode that injects a signal to continually track the phase difference (xy phase)

between the two orthogonal feeds. In addition, since the feeds are linearly polarized, there is very little contamination of the circular polarization signal by the total intensity signal. For further details on the ATCA instrumental polarization we refer readers to Sault, Rayner & Kesteven (2002).

A digital correlator (later replaced by the CABB broad-band digital correlator) allowed simultaneous observations in two bands, each with a bandwidth of 128 MHz divided into 32 channels. The observation frequencies within the 20 GHz band, covering the 16–24 GHz range, were chosen so as to maximize sensitivity, make use of optimal system temperatures and avoid correlator harmonics. The frequency bands were centred at 16.704 and 19.392 GHz. After calibration, data were averaged over the 256 MHz band so that our mean effective frequency was 18.048 GHz (hereafter tagged as ‘18’ GHz observations). The field of view (FOV) was ~ 2.6 arcmin.

The closeness of dishes in the H75 configuration can cause significant antenna shadowing for sources south of a declination of approximately -76° . Because the effects of shadowing and cross-talk on polarization measurements are unknown for the instrument, the sample for H75 observations was restricted to $-76^\circ < \delta < -30^\circ$. 11 $\delta < -76^\circ$ sources were observed for this project during a run that took place on 2006 October 17 with a more extended hybrid antenna configuration (H214) and observation bands centred at 18.752 and 21.056 GHz. Calibration and observation setup for these objects matched that of AT20G follow-up observations (see Murphy et al. 2010, for details). After the data reduction, the data for three of these objects (AT20G J115253–834410, AT20G J122454–831310, AT20G J155059–825807) at 20 GHz were found to be of poor quality in this run. However, they had good quality observations and a polarization detection in 2006 March during a previous AT20G follow-up run. Hence, we decided to include these measurements in the present analysis, flagging them in the catalogue.

For each run a bright point-like source was observed to calculate bandpass solutions and PKS 1934–638 was used as the primary flux density calibrator. The sample was broken into groups of 4–6 sources, located in the same sky region. Before each group was observed, an antenna pointing correction was performed using a nearby bright source to maintain directional accuracy despite the vast range of sky positions that the telescope had to observe in a short period of time. Each target was observed in two 70 s snapshots separated by 4 h.

Five targets were identified as extended within the primary beam, according to the extendedness criteria described in Murphy et al. (2010). As the flux density estimation methods that we use (see Section 3) are well suited for compact objects, the flux densities of extended sources are likely underestimated by an unknown amount and the polarized flux density is potentially wrong. For this reason, they were included in the list of extended sources for which we extracted the polarization information from the *WMAP* 9 yr co-added maps (see Section 4). In the end, after dropping 4 of the 9 very extended sources observed by Burke-Spolaor et al. (2009) and the 5 sources found to be extended, we are left with 180 sources for which we have good quality ATCA data.

Five antennas were available in our 16.704 GHz band, while there was an error in the 19.392 GHz band that left only four usable antennas at that frequency. To save observing time, no secondary calibrators were interleaved with the target observations, with the intention of self-calibrating each bright source during the data reduction. As described in Section 3, this turned out to be an unwise strategy in terms of polarization calibration. Nevertheless, 169 of the 180 extragalactic sources in our ATCA observed sample (94 per cent)

were detected in polarization at 18 GHz. However, only 145 of them also have $S_{18\text{ GHz}} > 500$ mJy in the 2006 October observations.

Table 2 provides an excerpt of the catalogue containing the data for the sample of 180 objects for which we have good quality ATCA data.

2.2 Lower frequency observations

Lower frequency observations were performed during a regular AT20G follow-up run (described in Murphy et al. 2010), carried out in 2006 November. In that epoch, we used a 5 antenna east–west array configuration, with a 1.5 km longest baseline. Two 30 s observations per source provided simultaneous 4800 and 8640 MHz snapshots. These observations had 128 MHz of bandwidth per frequency. The FOV were of 9.9 and 5.5 arcmin, respectively.

Massardi et al. (2011b) demonstrated that the median variability of total intensity flux densities for sources in flux density ranges similar to those of our sample over three month time-scale is 3.5 and 6.3 per cent, respectively, at ~ 5 and ~ 9 GHz, slightly increasing with time lag. Kurinsky et al. (2013) found a median variability index for total intensity flux densities of 1.0 per cent at 5 GHz in less than two months time lag for a similar sample in the Northern hemisphere. Hence, we could assume that our high- and low-frequency observation comparisons, on shorter time-scales, are not substantially affected by variability for most sources.

173 and 172 sources have good quality flux densities, respectively, at 4.8 and 8.6 GHz. 172 sources have good quality flux densities at all the frequencies (quality controls are described in Murphy et al. 2010). Of these, 137 have $S_{18\text{ GHz}} > 500$ mJy in the 2006 October run. 143 sources have a polarization detection at all the frequencies (79 per cent of the main sample), 119 of which have $S_{18\text{ GHz}} > 500$ mJy in the 2006 October run.

3 DATA REDUCTION

The 16.7 and 19.4 GHz data were reduced using the MIRIAD software package (Sault, Teuben & Wright 1995). The two frequencies were calibrated independently and then combined for 18 GHz imaging and flux density assessment. Opacity correction and a correction for the time-dependent instrumental χ phase difference was applied upon loading all data into MIRIAD. After this correction, a small residual gain offset still remained to be corrected in the following calibration stages. Bandpass solutions and primary flux density calibration were calculated and applied using PKS 1921–293 and PKS 1934–638, respectively.

For polarimetric calibration with calibrators of unknown polarization and sparse data (such as in our short snapshot observations), the standard MIRIAD procedure suggests calculating the largely stable instrumental leakage terms by using an unpolarized primary calibrator. The remaining polarization and gain terms are then calculated for each secondary phase calibrator.

Roughly 75 per cent of the sources in our sample are registered in the ATCA calibrator data base; all are sufficiently bright to determine adequate calibration solutions. However, though it is suggested that accurate Q and U values could be calculated from a relatively small amount of data, it was apparent that this was not the case for our ~ 3 min observations. Many solutions failed, while others produced impossible values for Q and U levels. This is likely due to the large number of free parameters and insufficient data length, even using the smallest possible solution interval. To overcome this hindrance, all polarization and gain solutions for the main sample were determined using the secondary calibrators that were interlaced in

Table 2. Catalogue of the first 10 sources of the extragalactic sample observed in the AT20G run dedicated to high-sensitivity polarization. The full source list, available as supplementary material online, includes 187 objects. Columns are: (1) AT20G name; (2) 18 GHz scalar average flux density (error given by equation 3); (3–5) polarized flux density, fractional polarization, and polarization angle at 18 GHz; (6–9) total intensity, polarized flux density, fractional polarization and polarization angle at 8.6 GHz; (10–13) total intensity and polarized flux densities, polarization degree and polarization angle at 4.8 GHz; (14) ratio between the visibilities amplitude averaged over the long baselines (~ 4.5 km) and the short ones (~ 0.2 km) in the AT20G catalogue (Chhetri et al. 2013); (15) flags: ‘s’ for sources with low declination observed during the AT20G 20 GHz run (i.e. the highest frequency is 20 GHz instead of 18 GHz), ‘f’ for sources with data at 20 GHz observed in 2006 March because of the poor quality of their data in the 2006 October runs, ‘b’ for very extended sources included in the Burke-Spolaor et al. (2009) sample, ‘w’ for extended sources which have had flux densities (or their upper limits) extracted from the WMAP 23 GHz 9-yr co-added maps (i.e. the highest frequency is 23 GHz instead of 18 GHz) and ‘c’ for sources candidate to be suitable polarization calibrators at millimetre and submillimetre wavelengths.

Name ‘AT20G’	S_{18} (mJy)	P_{18} (mJy)	Π_{18} (per cent)	ϕ_{18} ($^\circ$)	$S_{8.6}$ (mJy)	$P_{8.6}$ (mJy)	$\Pi_{8.6}$ (per cent)	$\phi_{8.6}$ ($^\circ$)	$S_{4.8}$ (mJy)	$P_{4.8}$ (mJy)	$\Pi_{4.8}$ (per cent)	$\phi_{4.8}$ ($^\circ$)	6 km Vis	Flags
J000435-473619	798.4 \pm 16.0	7.09 \pm 0.4	0.89 \pm 0.06	-78.70	929.0 \pm 0.7	17.60 \pm 0.8	1.90 \pm 0.01	-51.50	865.0 \pm 0.6	15.70 \pm 0.7	1.82 \pm 0.01	-37.04	0.98	...
J001035-302748	595.8 \pm 11.9	31.16 \pm 1.3	5.23 \pm 0.24	-46.66	676.0 \pm 0.7	25.40 \pm 0.9	3.75 \pm 0.01	-38.97	582.0 \pm 0.6	6.40 \pm 0.6	1.10 \pm 0.01	-52.22	0.96	c..
J001259-395426	1030.0 \pm 20.6	33.67 \pm 0.9	3.27 \pm 0.11	-86.42	1622.0 \pm 0.8	24.70 \pm 0.8	1.52 \pm 0.01	-85.56	1651.0 \pm 0.6	20.60 \pm 0.6	1.25 \pm 0.01	-88.13	0.96	...
J002616-351249	1136.0 \pm 22.8	1.65 \pm 0.4	0.15 \pm 0.04	-54.21	382.0 \pm 0.8	<2.40	133.0 \pm 0.7	<2.09	0.98	...
J004959-573827	1859.0 \pm 37.2	71.41 \pm 1.6	3.84 \pm 0.11	35.39	2460.0 \pm 0.8	67.10 \pm 0.9	2.73 \pm 0.01	9.14	2237.0 \pm 0.7	67.00 \pm 0.7	2.99 \pm 0.01	5.78	0.97	c..
J005109-422632	683.7 \pm 13.7	11.11 \pm 0.5	1.62 \pm 0.08	-2.21	950.0 \pm 0.7	10.00 \pm 0.7	1.31 \pm 0.07	8.64	1060.0 \pm 0.6	20.00 \pm 0.6	1.46 \pm 0.06	10.43
J005846-565911	785.1 \pm 15.7	12.60 \pm 0.6	1.60 \pm 0.09	-23.19	709.0 \pm 0.6	16.70 \pm 0.6	2.36 \pm 0.01	-72.10	573.0 \pm 0.6	<1.80	0.99	...
J010645-403419	3785.0 \pm 75.7	78.84 \pm 2.9	2.08 \pm 0.09	-8.26	3680.0 \pm 1.5	60.00 \pm 1.5	1.70 \pm 0.04	46.43	2410.0 \pm 0.7	30.00 \pm 0.7	1.32 \pm 0.03	-62.00
J010915-604948	375.0 \pm 7.5	15.73 \pm 0.6	4.20 \pm 0.18	-14.55	516.0 \pm 0.7	7.40 \pm 0.7	1.44 \pm 0.01	-49.53	533.0 \pm 0.6	9.10 \pm 0.6	1.70 \pm 0.01	88.12	0.97	...
J012457-511316	369.5 \pm 7.4	7.00 \pm 0.5	1.90 \pm 0.13	-11.21	298.0 \pm 0.7	9.60 \pm 0.7	3.20 \pm 0.01	-15.13	220.0 \pm 0.7	2.40 \pm 0.7	1.11 \pm 0.01	-31.11	0.99	...

our observations, though these were originally intended for use with the objects presented by Burke-Spolaor et al. (2009). Merged solutions from the eight available calibrators afforded an observation every 1–2 h in various regions of the sky. Because each was typically observed at least for 1 min at any of eight parallactic angle intervals spanning approximately 6 h, they had sufficient parallactic angle coverage to disentangle instrumental polarization from the intrinsic calibrator Stokes Q and U levels. The absolute flux density scale calculated from PKS 1934–638 was then used to bootstrap the secondary calibrators; gain and instrumental polarization terms were then applied to all target sources.

Incidentally, tests run on the calibrated sources to check levels of residual phase instability, which is usually due to imperfect phase calibration (as given in percentage by dividing the source vector amplitude by its scalar amplitude) showed that using calibrators observed frequently in time can give a better phase calibration than less frequently observed calibrators even if they are closer in space to the target source (as shown by 0–20 per cent residual decorrelation in our sample versus the 0–50 per cent found in AT20G data over similar time-scales and weather conditions).

Hence, in polarimetry experiments covering large areas of the sky, it appears more pertinent to have many observations of one calibrator throughout an observation and therefore have sufficient data to determine accurate polarization solutions, despite possible non-proximity to target sources. However, calibrators in this experiment were never further than 30° from any target source and were restricted to high-elevation observations.

The 4.8 and 8.6 GHz data for the main sample and the 20 GHz data for the 11 $\delta < -76^\circ$ sources had observational modes exactly coincident with the AT20G survey follow-up, and thus were flagged and reduced using the automatic pipeline developed for the AT20G (Murphy et al. 2010).

Stokes I intensities were determined from the visibilities to avoid the inclusion of phase instabilities inherent in image-based measurements. This technique takes the scalar average of the visibility amplitudes and is robust for bright (>200 mJy), point-like sources only.

To acquire Stokes Q , U and V flux densities, images were created and deconvolved using the MIRIAD task CLEAN. To correct the Stokes Q , U and V images for decorrelation, we took advantage of the fact that Stokes parameters, simultaneously measured, are affected by decorrelation originating in atmospheric phase instabilities (as might be left after imperfect calibration). We can thus use the fractional level of residual decorrelation (χ) in Stokes I , calculated and applied to Q , U , and V flux densities as

$$\chi = \frac{I_{\text{sca}}}{I_{\text{map}}}, \quad (1)$$

$$Z = \chi \cdot Z_{\text{map}}, \quad (2)$$

where Z represents Stokes Q , U or V , I_{sca} is the scalar-averaged Stokes I flux density, and I_{map} and Z_{map} represent the values at the position of the peak Stokes I emission in the relevant image. The image peak for all sources was sufficient to determine the decorrelated flux density measurements; the pixel size was typically 10 arcsec, and no sources in this subsample were extended significantly beyond this.

The polarized intensity and the position angles were then calculated using standard first-order debiasing, where $P = \sqrt{Q^2 + U^2 - \sigma_V^2}$ (Wardle & Kronberg 1974; Simmons & Stewart 1985). The last term, σ_V , is the rms noise in the Stokes V image. Most extragalactic sources do not have significant levels of circular

polarization; therefore, the Stokes V signal is usually undetected or very low for all such sources. This point makes the noise level in Stokes V a reasonable estimate of the background noise level in the Q and U images, and thus gives a good estimate for the debiasing correction.

The polarization angles are given by $\phi = (1/2) \arctan(U/Q)$ and the linear polarization degree, Π , is given in the percentage: $\Pi = 100 \cdot P/I$.

3.1 Error budget

The rms scatter σ_V provided a measurement of the noise on the scalar average flux density of the order of 1–2 per cent. Telescope pointing inaccuracies, considering the ~ 15 arcsec pointing errors and the 18 GHz primary beam response function are expected to cause a possible attenuation of up to 2 per cent in all Stokes parameters.

Errors in the primary flux density scaling from PKS 1934–638 are estimated by comparing the online ATCA calibrator catalogue data at 18.496 GHz to the high-frequency polynomial model used by MIRIAD to calculate the scaling factor. The model predicts $I_{\text{model}} = 1.0259$ Jy at our observing frequency while the average of values measured in 2006 is $I_{\text{data}} = 1.0278$ Jy; the calibration error in source flux densities is thus approximately 0.2 per cent. Given the much larger statistical errors in I_{sca} , we ignore this calibration error. The net error in total intensity is thus given by

$$\sigma_I^2 = (0.02 I_{\text{sca}})^2 + \sigma_V^2. \quad (3)$$

Errors in the Stokes parameters Q , U and V can arise from leakage of the much brighter Stokes I signal due to imperfections in the alignment of orthogonal receiver components. The correction of this effect using PKS 1934–634 and an iterative leakage calculation using secondary calibrators results in a negligible error compared to the system noise. The noise term is calculated by propagation of errors through equations (1) and (2). The main contributions to the global error come from the antenna pointing inaccuracy and from the noise estimated by the rms levels in off-source regions in the restored image

$$\sigma_Z^2 = \left(\frac{I_{\text{sca}} \sigma_{Z_{\text{map}}}}{I_{\text{map}}} \right)^2 + \left(\frac{Z I_{\text{sca}} \sigma_{I_{\text{map}}}}{I_{\text{map}}} \right)^2 + (0.02Z)^2, \quad (4)$$

where Z is either Q or U . The error on the polarized intensity can then be derived as

$$\sigma_P^2 = \frac{Q^2 \sigma_Q^2 + U^2 \sigma_U^2}{Q^2 + U^2}, \quad (5)$$

where Q and U are calculated from equation (2). Note that if $\sigma_Q \simeq \sigma_U$ (as expected for low polarization sources in noisy maps), then $\sigma_P = \sigma_Q = \sigma_U$. We defined as a non-detection of a source $P < 3\sigma_P$, and used $3\sigma_P$ as the upper limit on the polarized flux density for such sources. The errors in the fractional polarization were obtained with the usual error propagation from σ_P and σ_I .

4 POLARIZATION OF EXTENDED SOURCES

As mentioned above, our ATCA data on extended sources are largely incomplete. Of the nine very extended sources with $S_{20\text{GHz}} > 0.5$ Jy selected by Burke-Spolaor et al. (2009) for wide field imaging and polarimetry, one (Fornax A) was not observed due to its highly diffuse emission, most of which would remain undetected even in ATCA's most compact configuration, and for three more (AT20G J013357–362935, AT20G J132527–430104 and AT20G

Table 3. Polarized and total flux densities measurements (in mJy) for the three extended sources with a polarization detection in the *WMAP* 9 year co-added full-sky maps.

Name	$P_{23\text{ GHz}}$	$P_{33\text{ GHz}}$	$P_{41\text{ GHz}}$	$P_{64\text{ GHz}}$	$S_{23\text{ GHz}}$	$S_{33\text{ GHz}}$	$S_{41\text{ GHz}}$	$S_{64\text{ GHz}}$
Fornax A (RA:03:22:41.7; Dec.:−37:12:30)	1074 31	867 44	589 64	<354	9321 134	5350 184	3275 173	905 255
PicA – AT20GJ051949–454643	457 35	372 50	280 82	484 137	6464 207	5661 235	4656 210	3139 270
CenA – AT20GJ132527–430104	3322 70	2699 81	2323 120	2075 173	51 006 260	41 909 248	35 731 245	26 767 335

J133639–335756) only subregions containing compact structure were observed. Moreover, five objects in our initial sample were found to be extended with respect to our synthesized beam and therefore to have unreliable polarized flux density estimates in the AT20G catalogues.

For the sake of completeness, we have attempted to estimate the polarized flux densities of these extended objects using the *WMAP* 23 GHz 9 yr co-added map, where the $0^{\circ}88$ beam is collecting most of the extended emission seen by ATCA.

The extraction of the total intensity flux density was performed using the IFCAMEX software package,¹ which has been used in the past to extract flux densities from *WMAP* (Massardi et al. 2009) and *Planck* data (Planck Collaboration 2013 results XXVIII). The extraction of polarized flux densities from the *WMAP* 9 yr data has been performed using the IFCAPOL software package used to characterize polarized sources in *WMAP* 5 yr maps (López-Caniego et al. 2009). This software implements the Filtered Fusion approach (Argüeso, Sanz & Herranz 2011), where a maximum likelihood estimator is obtained for the Q and U maps of each source. As a result, de-noised Q_f and U_f maps are produced and the polarized flux density at the position of the source is obtained from the map of $P = \sqrt{Q_f^2 + U_f^2}$.

Note that the *WMAP* polarization maps are very noisy and it is important to assess whether or not our estimate of the polarized flux density at the position of a source detected in the total intensity comes from the source or from a maximum of the CMB at that position. This is done by assessing the significance of each detection/estimation in the P map. For each source, we calculate the 99.90 per cent significance level, as explained in López-Caniego et al. (2009), and check that our estimate of the flux density at the position of the source is at least above this level. This allows us to discriminate between truly significant detections in the maps of P from random peaks of the background.

Suitable detections were obtained at 23 GHz for 2 of the 9 sources in the above mentioned subsample (Fornax A and Centaurus A). Extractions allowed us to define an upper limit of the integrated polarized flux density for the other seven cases, but for two of these objects the extraction algorithm could not determine the total flux density.

The results of this extraction are listed in the main catalogue (see Table 2), flagged with ‘w’. The 23 GHz flux densities for the extended sources will be included in the following analysis without any correction for the spectral behaviour between the *WMAP* observing frequency and the 18 GHz band. Hence, the full sample that will be used in the next section includes 187 sources; it is 99 per cent complete with $S_{20\text{ GHz}} > 500$ mJy at the 2006 survey selection epoch. The polarized emission detection rate is 91.4 per cent.

Fornax A is one of the closest and most extended sources in the Southern hemisphere, with two lobes extended over a region more than 50 arcmin wide. Its weak core has a flux density at 20 GHz

much below the AT20G survey 10 mJy detection limit, and for this reason it is not included in the AT20G catalogue. *WMAP* was able to integrate over the source area only at 23 GHz, which we included in our analysis.

In addition, a detection of polarized flux density at 23 GHz in the *WMAP* maps has also been obtained for Pictor A. It is among the targets observed in mosaic mode by Burke-Spolaor et al. (2009), where it was identified as the best polarization calibrator among the extragalactic sources for arcminute resolution experiments (and in particular for the *Planck* experiment), despite its steep spectrum in the region ~ 1 –20 GHz. The detected value of 457 ± 35 mJy at 23 GHz (listed in Table 3) is comparable with the 500 ± 60 mJy measured for polarized flux density at 18 GHz by Burke-Spolaor et al. (2009) over the whole source (listed in Table 2 and used in the following analysis). The *WMAP* detections seem to indicate a steep spectrum in polarized emission in the *WMAP* frequency range.

Table 3 lists the flux densities in total intensity and polarization for ForA, PicA and CenA in all the *WMAP* bands. Notice that the detections at frequencies above 23 GHz might refer only to fractions of the sources if they are more extended than the *WMAP* beams.

5 DATA ANALYSIS

5.1 Spectral properties of the sample

We have defined the spectral index α as $S \propto \nu^\alpha$. The analysis of AT20G data by Chhetri et al. (2012) has confirmed that $\alpha = -0.5$ is a physically meaningful threshold to separate compact, self-absorbed sources from structurally complex, extended objects. The majority (105, i.e. 61 per cent) of our sources are flat spectrum objects (‘F’ $-0.5 < \alpha_{4.8}^{8.6}, \alpha_{8.6}^{18} < 0.5$). The remaining can be classified as follows: 25 (15 per cent) sources have peaked spectra (‘P’, $\alpha_{4.8}^{8.6} > 0$ and $\alpha_{8.6}^{18} < 0$); 29 (17 per cent) have steep spectra (‘S’, $\alpha_{4.8}^{8.6} < 0$ and $\alpha_{8.6}^{18} < 0$); 12 (7 per cent) have inverted (‘I’, $\alpha_{4.8}^{8.6} > 0$ and $\alpha_{8.6}^{18} > 0$); only 1 source has upturning spectra (‘U’, $\alpha_{8.6}^{18} > 0$). Fig. 1 illustrates the spectral classification of the 172 sources in the sample with total intensity data at all the three frequencies (4.8, 8.6 and 18 GHz). The different distribution of the sources in the ($\alpha_{4.8}^{8.6}, \alpha_{8.6}^{18} < 0.5$) plane indicates that the spectral properties over the 5 to 18 GHz frequency range may be different in total intensity and in polarization (see also Fig. 2). The median values of the 4.8–8.6 GHz and 8.6–18 GHz spectral indices in polarization are 0.20 and -0.16 , respectively. This effect could be a combination of Faraday depolarization operating at the lower frequencies, superposition of multiple components with different polarized spectra and different magnetic field properties for the components that dominate the emission at the different frequencies.

We constructed the matrix in Table 4 by comparing the total intensity and the polarization spectral behaviours. The distribution across the cells confirms the differences of spectral behaviour in polarization and in total intensity. Similar behaviours translate in a diagonal matrix. However, even if a tiny effect due to Faraday depolarization could be the cause of the high number of peaked spectra in

¹ <http://max.ifca.unican.es/IFCAMEX>

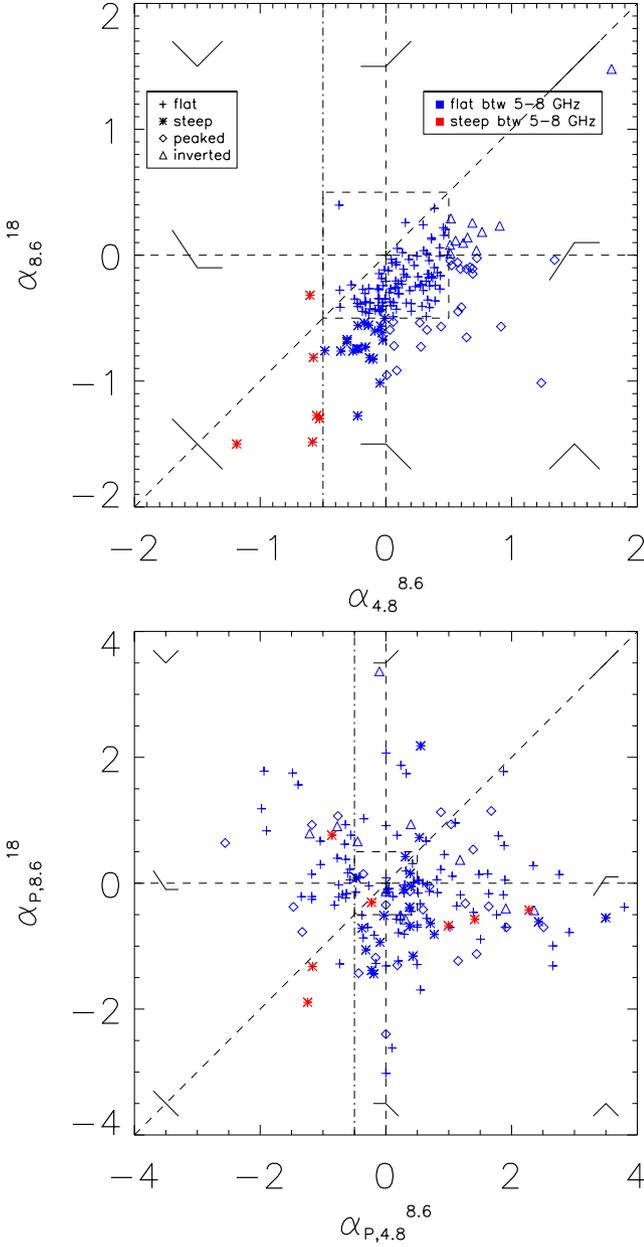


Figure 1. Radio colour–colour diagram for (from top to bottom) flux density and polarized flux density. Error bars and upper limits have been omitted for clarity of display.

polarized emission (because of the lower level of emission at lower frequencies), more than ~ 30 per cent of sources show a polarized emission spectral index on the lower frequency range higher than at the higher frequencies, indicating that Faraday emission is not significantly affecting their spectral behaviour.

Chhetri et al. (2012) demonstrate that low-frequency spectral index selections in flat and steep populations are more effective in identifying compact and extended objects than high-frequency spectral indices. For this reason, we rely on $\alpha_{4.8}^{8.6}$ to classify flat and steep spectra sources in some of the following analysis. Fig. 1 also shows that for this selection criterion the total intensity and polarization spectral behaviours are different. The median of spectral indices in total intensity and polarized flux densities for each

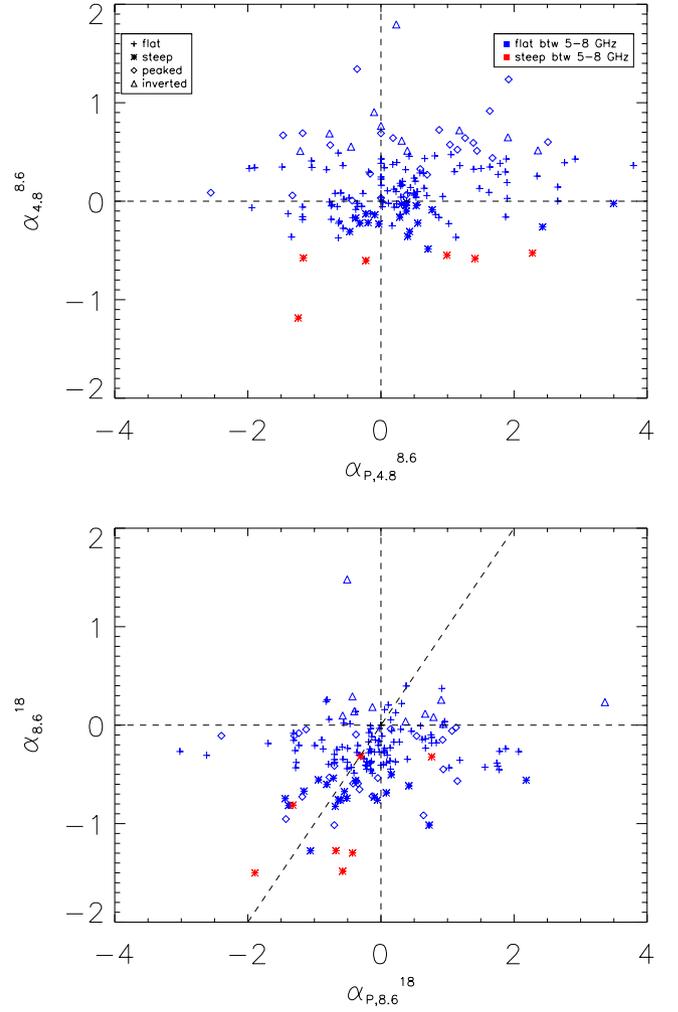


Figure 2. Comparison between the spectral index in total intensity and in polarization in the ranges 4.8–8.6 GHz (upper panel) and 8.6–18 GHz.

Table 4. Matrix of the number of sources classified according to both the total intensity and polarization spectral behaviour. The columns are the spectral shape in polarization, the rows the spectral shape in total intensity. The spectral types are defined in the text.

S \rightarrow	(I)	(P)	(F)	(S)	(U)
Pol. \downarrow					
Inverted (I)	2	4	21	2	0
Peaked (P)	5	10	33	10	0
Flat (F)	1	3	17	7	1
Steep (S)	0	4	15	9	0
Upturning (U)	4	3	17	1	0

class and for the full sample are in Table 5. An overall steepening is appreciable both in total intensity and in polarized emission.

5.2 Fractional polarization

We derived the distribution of the polarization fractions in Fig. 3 and Table 6 using a bootstrap re-sampling method in order to account for the uncertainties in both the polarized and the total intensity flux density measurements.

Table 5. Median values of spectral indices in different frequency ranges in total intensity and polarization. Spectral classes are selected according to their behaviour between 4.8 and 8.6 GHz.

		4.8–8.6 GHz	8.6–18 GHz	4.8–18 GHz
α_{v1}^{v2}	All	0.09	−0.28	−0.11
	Flat	0.12	−0.27	−0.08
	Steep	−0.60	−1.27	−0.95
$\alpha_{p,v1}^{v2}$	All	0.20	−0.16	−0.006
	Flat	0.23	−0.15	−0.006
	Steep	−0.23	−0.43	0.04

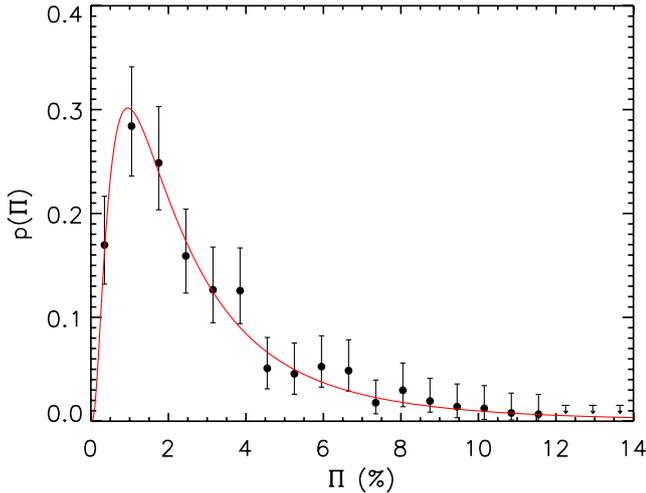


Figure 3. Distributions of the polarization degree at 18 GHz. Errors and upper limits correspond to a 68 per cent c.i. The solid line shows the log-normal distribution with median fractional polarization 2.14 per cent and $\sigma = 0.90$ per cent (see the text for details).

Table 6. Distributions of the fractional polarization at 18 GHz for the full sample.

Π (per cent)	Probability	−	+
0.35	0.170	0.038	0.047
1.05	0.284	0.048	0.057
1.75	0.249	0.045	0.054
2.45	0.159	0.036	0.045
3.15	0.127	0.032	0.041
3.85	0.126	0.032	0.041
4.55	0.051	0.020	0.030
5.25	0.046	0.020	0.030
5.95	0.053	0.020	0.030
6.65	0.049	0.020	0.030
7.35	0.018	0.011	0.022
8.05	0.030	0.016	0.026
8.75	0.019	0.011	0.021
9.45	0.014	0.011	0.021
10.15	0.012	0.011	0.021
10.85	0.008	0.007	0.019
11.55	0.007	0.007	0.019
12.25	<0.015		
12.95	<0.015		
13.65	<0.015		
14.35	<0.015		

We generated 1000 simulated catalogues by re-sampling, with repetitions, the input catalogue of polarized flux densities. In each simulation, values for the polarized and the total intensity flux densities were randomly assigned to each source by assuming a Gaussian distribution with a mean equal to the measured values and a σ equal to the quoted errors. When only an upper limit was available on the polarized flux density, we generated random values between 0 and the quoted upper limit assuming a uniform distribution. For each realization, the polarization fraction is estimated for each object as the ratio between the simulated polarized flux density and the simulated total intensity flux density. The resulting values are then distributed into bins of polarization fraction.

The final distribution of the polarization fractions is given by the mean value of the simulated polarization fractions in each bin with uncertainties derived assuming a Poisson statistic, according to the prescriptions of Gehrels (1986). In Fig. 3, error bars correspond to the 68 per cent confidence interval.

The solid line is a fit assuming the log-normal distribution

$$f(\Pi) = \text{cost} \cdot \frac{1}{\sqrt{2\pi}\sigma\Pi} e^{(-0.5(\ln(\Pi/\Pi_m))^2/\sigma^2)},$$

with $\text{cost} = 0.97$ and $\sigma = 0.90$ per cent; $\Pi_m = 2.14$ per cent is the median value of the distribution.

Table 7 summarizes the mean and quartiles of the distributions of fractional polarization at 4.8, 8.6 and 18 GHz as calculated with survival analysis techniques to hold for upper limits in polarized flux density for each of the spectral classes identified considering the spectral behaviour in the ranges 4.8–8.6 and 8.6–18 GHz (see Fig. 1) as discussed in the previous section. While no significant trend is visible with frequency for flat, upturning or peaked spectra objects, there is a tiny indication that the mean polarization fraction for steep spectrum sources decreases as the frequency decreases, probably under the effect of Faraday depolarization, but the significance is poor because we deal with fairly broad distributions.

Table 8 shows the parameters describing the distributions of fractional polarization for the full sample and for the flat and steep spectrum objects as calculated with survival analysis techniques to hold for upper limits in polarized flux density, selected according to the spectral behaviour in the 4.8–8.6 GHz frequency range only. There is a general trend towards an increase of the mean fractional polarization as the frequency increases for the whole sample and for each spectral class (even if we have poor statistics for the steep-spectrum sources). This trend is not so noticeable if we consider the median values, similarly to the findings in Table 7, as we deal with fairly broad distributions (see Table 8).

In fact, the differences among the polarization degree distributions at different frequencies are not statistically significant. All the generalized Wilcoxon two-sample tests available in the `ASURV` package reject the null hypothesis, which states that the distributions of fractional polarization are drawn from the same parent distribution, for the 4.8 and 8.6 GHz distributions for the full sample. In the range 8.6–18 GHz a correlation is found by the different methods at $\sim 2\sigma$ – 3σ significance level. Between 4.8 and 18 GHz the Gehan's, and Peto and Peto version of the Wilcoxon test and the log rank test confirm a correlation at the $\sim 2\sigma$ significance level, while the Peto and Prentice version indicate a $\sim 4\sigma$ level (see Fig. 4).

We then conclude that, for our sources, there is no statistically significant evidence of an increase of the median polarization degree with a frequency above 4.8 GHz. This implies on the one hand that already at 4.8 GHz the Faraday depolarization is not very important, and on the other hand that the magnetic field is not substantially

Table 7. Parameters describing the distributions of fractional polarization at 4.8, 8.6 and 18 GHz for each spectral class defined according to Fig. 1. The spectral types are defined in the text (upturning class has been neglected because it includes only one source). For each frequency and spectral class we quote the number of detections, the mean fractional polarization and its error, the first, second and third quartiles of the distribution.

	F	I	P	S
18 GHz				
N_{TOT}	105	12	25	29
Detections	101	11	25	25
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.57 ± 0.82	1.55 ± 0.25	3.24 ± 0.53	3.49 ± 0.59
1 2 3 quartiles	1.21 2.06 3.68	1.00 1.37 2.08	1.13 2.11 4.78	0.25 2.72 6.02
8.6 GHz				
N_{TOT}	105	12	25	29
Detections	100	10	24	24
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.35 ± 0.16	1.28 ± 0.29	2.64 ± 0.42	2.85 ± 0.51
1 2 3 quartiles	1.10 2.06 3.18	0.30 1.23 1.70	1.07 1.85 3.67	0.17 1.62 4.89
4.8 GHz				
N_{TOT}	105	12	25	29
Detections	93	10	24	21
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.13 ± 0.15	1.48 ± 0.30	2.59 ± 0.31	2.19 ± 0.46
1 2 3 quartiles	0.95 1.85 2.98	0.55 1.32 1.95	1.40 2.47 3.12	0.00 1.29 3.69

Table 8. Parameters describing the distributions of fractional polarization for the full sample and for the flat ($\alpha_{\nu 1}^{\nu 2} > -0.5$) and steep ($\alpha_{\nu 1}^{\nu 2} < -0.5$) spectrum objects selected in the frequency ranges $\nu 1 = 4.8$ and $\nu 2 = 8.6$ GHz. For each frequency and spectral class we quote the number of detections, the mean fractional polarization and its error, the first, second and third quartiles of the distribution, and the probability that flat and steep spectrum objects are drawn from the same parent distribution, according to the two-sample Wilcoxon test.

	Full sample	Flat selected between 4.8 and 8.6 GHz	Steep
18 GHz			
N_{TOT}	187	163	9
Detections	171	157	5
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.79 ± 0.17	2.76 ± 0.17	2.32 ± 0.99
1 2 3 quartiles	1.09 2.04 3.84	1.20 2.08 3.82	0.01 0.55 2.54
Prob (flat–steep)	7.8 per cent		
8.6 GHz			
N_{TOT}	172	163	9
Detections	158	151	7
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.38 ± 0.15	2.42 ± 0.15	1.80 ± 0.77
1 2 3 quartiles	0.94 1.80 3.27	1.03 1.91 3.28	0.04 0.55 1.46
Prob (flat–steep)	5.3 per cent		
4.8 GHz			
N_{TOT}	173	164	9
Detections	149	143	6
$\langle \Pi \rangle \pm \sigma_{\langle \Pi \rangle}$	2.16 ± 0.13	2.21 ± 0.13	1.08 ± 0.54
1 2 3 quartiles	0.87 1.82 3.05	0.94 1.91 3.07	0.00 0.26 1.22
Prob (flat–steep)	0.1 per cent		

more ordered in the regions dominating the emission at higher frequencies.

This conclusion is in line with the results by Battye et al. (2011) who found that the fractional polarization of their sources, selected from the *WMAP* point source catalogue and mostly flat spectrum, is almost independent of frequency in the range 8.4–43 GHz, with

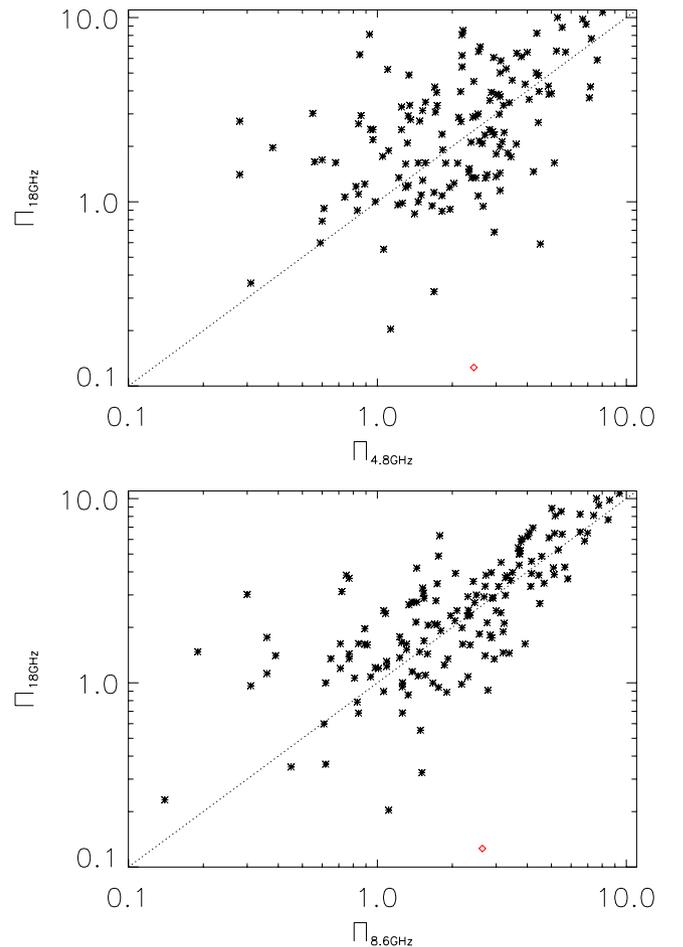


Figure 4. Fractional polarization at 18 GHz as a function of 4.8 and 8.6 GHz for the full sample (red diamonds indicate upper limits at one or both the frequencies).

median values in the range 2–2.5 per cent. Klein et al. (2003) also found that flat spectrum sources in their sample selected at 408 MHz are characterized by almost constant polarization degrees, with median values $\simeq 2.5$ per cent, over the frequency range 2.7–10.5 GHz. On the other hand, they reported a steady increase of the polarization degree of steep-spectrum sources from 2.2 per cent at 1.4 GHz to 5.8 per cent at 10.5 GHz. This increase, however, can be largely due to the bias induced by the requirement of polarization detection at high frequencies: since steep-spectrum sources become faint at high frequencies, only those with exceptionally high polarization degrees are detected. The same bias may explain the increase with frequency of the median polarization degree observed by Sajina et al. (2011) for the subset of their AT20G sources with detected polarized flux in all four bands (4.86, 8.46, 22.46 and 43.34 GHz). This interpretation is consistent with their finding that the trend is stronger for steeper spectrum sources as well as for the lower flux density sources.

Several studies of radio source polarization, mostly for samples collected at 1.4 GHz, dominated by steep-spectrum sources, have reported indications that the polarization degree increases with decreasing flux density (Mesa et al. 2002; Tucci et al. 2004; Taylor et al. 2007; Grant et al. 2010; Subrahmanyan et al. 2010). A similar trend was also found by Sadler et al. (2006) for AT20G sources with $S_{20\text{GHz}} > 100$ mJy. It was, however, not confirmed by the analysis of the full AT20G sample by Massardi et al. (2011a), although the relatively low detection rate in polarization prevented a clear conclusion. Our data, with a high detection rate, do not show any significant trend of Π with flux density at any of the frequencies (4.8, 8.6 and 18 GHz) for $S_{18\text{GHz}} > 500$ mJy, with only 45 per cent probability the null hypothesis that sources with flux densities above and below 1 Jy at 18 GHz come from the same parent distribution is true. It must, however, be noted that our flux density range is limited and substantially narrower than that of Sadler et al. (2006).

The median of the ratio between the fractional polarization at two different frequencies provides an estimate of the depolarization as the frequency decreases. We found that $\langle \Pi_{4.8\text{GHz}} / \Pi_{18\text{GHz}} \rangle = 79.7$ per cent and $\langle \Pi_{8.6\text{GHz}} / \Pi_{18\text{GHz}} \rangle = 86.9$ per cent, which implies a median depolarization of 20.3 and 13.1 per cent, respectively, in the two ranges of frequency 4.8–18 and 8.6–18 GHz. These findings confirm the results of Massardi et al. (2011a) and those by Tucci & Toffolatti (2012).

6 SOURCE COUNTS IN POLARIZATION

We derived the differential number counts in the polarized flux density in two ways. Results are listed in Table 9 and shown in Fig. 5. In both cases, we used the bootstrap resampling method and performed 1000 simulations.

As a first method, we directly measured the number counts from the catalogue of polarized flux density measurements. For each simulation, the catalogue was resampled with repetitions. As before, a value for each of the polarized flux density was randomly assigned to each source by assuming a Gaussian distribution with a mean equal to the measured values and σ equal to the quoted errors. When only an upper limit was available on the polarized flux density, we generated random values between 0 and the quoted upper limit assuming a uniform distribution. The distribution of the polarized flux densities derived from the simulations were binned into a histogram, and the mean value in each bin was taken as the measurement of the number count in that bin, after dividing by the bin size and by the survey area. The errors on the number counts were derived assuming a Poisson statistic, according to the prescriptions

Table 9. Source counts at 18 GHz (including also *WMAP* data) as estimated for the current sample (upper panel) and for the whole AT20G sample (lower panel). See the text for details. Errors enclose the range of 68 per cent c.i.

Sample	$\log [P(\text{Jy})]$	$dN/d\log P$ (deg $^{-2}$)	–	+	
This paper	– 1.428	0.0119	0.0026	0.0028	
	– 1.238	0.0086	0.0020	0.0027	
	– 1.047	0.0042	0.0015	0.0019	
	– 0.857	0.0013	0.0006	0.0017	
	– 0.666	0.0014	0.0007	0.0016	
	– 0.476	0.0005	0.0004	0.0012	
	– 0.285	0.0005	0.0004	0.0012	
	– 0.095	<0.0009			
	0.096	0.0005	0.0004	0.0012	
	0.287	<0.0009			
	0.477	0.0005	0.0004	0.0012	
	Whole AT20G	– 2.000	0.1115	0.0055	0.0053
		– 1.809	0.0624	0.0040	0.0041
– 1.619		0.0353	0.0030	0.0031	
– 1.428		0.0203	0.0024	0.0025	
– 1.238		0.0114	0.0017	0.0021	
– 1.047		0.0062	0.0012	0.0017	
– 0.857		0.0032	0.0010	0.0011	
– 0.666		0.0016	0.0006	0.0009	
– 0.476		0.0008	0.0004	0.0008	
– 0.285		0.0003	0.0003	0.0005	
– 0.095		<0.0004			
0.096		<0.0004			
0.287		<0.0004			
0.477	<0.0004				

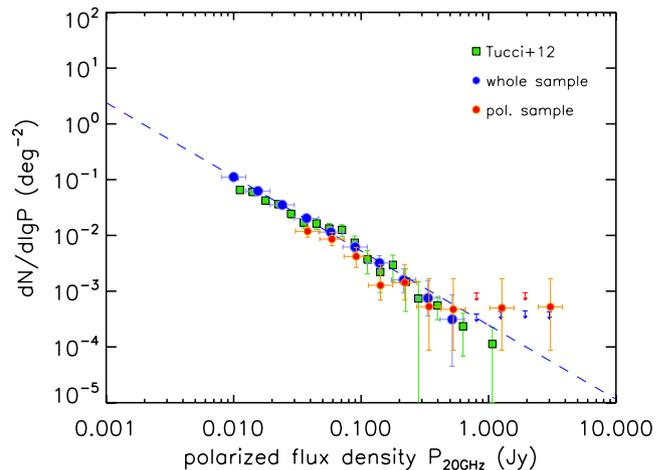


Figure 5. 18 GHz differential source counts in polarization calculated with the two methods described in the text compared with the findings by Tucci & Toffolatti (2012). The dashed line represents a linear fit of the whole AT20G sample data.

of Gehrels (1986). The results are shown in Fig. 5 by the red dots with error bars corresponding to the 68 per cent confidence interval.

For the second method, we started off with the Massardi et al. (2011a) catalogue of total intensity flux densities. In each simulation, the total intensity flux densities was randomly assigned to each source by assuming a Gaussian distribution with a mean equal to the measured values and sigma equal to the quoted errors. At the same time a realization of the polarization fraction was produced from the catalogue of polarized flux densities following the

procedure described before. The simulated total intensity flux density and the simulated value of the polarization fraction were then used to derive the corresponding polarized flux density of each object. Finally, the distribution of the polarized flux densities obtained from the simulations were binned into a histogram and the mean value in each bin was taken as the measurement of the number count in that bin, divided by the bin size and by the survey area. The errors on the number counts were derived assuming a Poisson statistic, according to the prescriptions of Gehrels (1986). The results are shown in Fig. 5 by the blue dots with error bars corresponding to the 68 per cent confidence interval. These points are linearly fitted with the relation

$$\log(dN/d \log P[\text{deg}^{-2}]) = (-1.32 \pm 0.05) \log P[\text{Jy}] \\ + (-3.60 \pm 0.09).$$

In the same figure the green squares represent the number counts in the polarized flux densities derived by Tucci & Toffolatti (2012).

7 POLARIZATION CALIBRATORS AT MILLIMETRIC WAVELENGTHS

Finding suitable calibrators for polarization studies at millimetric wavelengths is a difficult exercise. On the one hand, extragalactic sources typically have a very low level of circular polarization, which simplifies the calibration solutions. On the other hand, as demonstrated in the previous sections, the fractional polarization is typically a small fraction of the total intensity and polarized behaviour cannot be easily predicted on the basis of total intensity properties. Furthermore, source variability makes the fractional polarization unknown. Source compactness (with respect to the beam of the used telescope) allows us to consider only on-axis effects. Models of extended emissions of calibrators could be used to cope with off-axis effects, if there is no better suited compact calibrator.

Polarized sources allow us to recover gain and instrumental polarization parameters calibration, even if the polarization fraction of the calibrator is unknown (but non-zero) and its polarization angle is also unknown, by observing over a wide range of parallactic angles (for alt-az telescope mounts and assuming the instrumental parameters do not vary in the time of the observation; see Sault, Hamaker & Bregman 1996).

Statistical analysis clearly showed that spectral behaviours in total intensity and polarization are typically different, with a tiny indication that the fractional polarization increases with frequency, at least for high-frequency-selected steep-spectra sources. Therefore, it is difficult to make predictions with our data at frequencies $\gtrsim 20$ GHz. Any candidate selection requires proper monitoring at the observing frequency to confirm the calibrator properties. Hence, we stress the fact that in this section we do not claim any definition of effective criteria to identify good calibrators, but we only suggest a short list of targets for future calibrator monitoring programmes at frequencies $\gtrsim 20$ GHz.

The list is selected in the catalogue presented in the previous sections of sources

- (i) south of -30° , excluding the Galactic plane region ($|b| \leq 1.5^\circ$) and the LMC region
- (ii) with total intensity flux density at 20 GHz $S_{20\text{GHz}} > 500$ mJy in the 2006 AT20G selection.

Among them we selected the sources with fractional polarization at 20 GHz $\Pi_{20\text{GHz}} > 3$ per cent (corresponding to $P > 15$ mJy and, according to the distribution in Section 5, enclosing about

40 per cent of our sources) and at least two of the following properties

- (i) flat spectral index between 8 and 20 GHz $\alpha_8^{20} > -0.5$ in the 2006 epoch, to enhance the chances of choosing sources that are bright above 20 GHz;
- (ii) $S_{20\text{GHz}} > 500$ mJy in all the AT20G epochs, and varying less than 10 per cent with respect to the 2006 October epoch over 1 yr of observations or smaller than 20 per cent over longer epochs to select the most stable objects according to the available data;
- (iii) increasing fractional polarization between 8 and 20 GHz, to favour high levels of fractional polarization observations at the higher frequencies.

The selected sample contains 29 sources. They have been flagged in Table 2 with a ‘c’ and some of their properties are summarized in Table 10. All the sources are classified as point-like in the AT20G catalogue (i.e. with most of their emission within the synthesized beam of 10 arcsec at 20 GHz).

Furthermore, where available, we considered the source ‘6 km visibility’ which is the ratio of the scalar-averaged amplitudes at the long baselines (~ 4.5 km) over the short baselines (~ 0.2 km) of the AT20G catalogue. Chhetri et al. (2013) provided an efficient tool to select point-like sources in the AT20G, where 6 km visibility values larger than 0.86 at 20 GHz identify sources with angular size smaller than 0.15 arcsec, which makes them good calibrator candidates. Only two sources (AT20G J142432–491349 and AT20G J170918–352520) have ratios smaller than the threshold, but they could still be candidates for observations with synthesized beams on the scale of $\gtrsim 1$ arcsec.

As expected, spectral behaviour are mostly slowly steepening or peaked. We also note that Chhetri et al. (2012) found that compact sources show spectral steepening at rest frequencies ~ 30 GHz, which may affect the higher frequency spectra. Optical identifications (Mahony et al. 2011) indicate that the sample is mostly composed by QSO at mean redshift above 1.2. Almost all the objects are calibrators in the ATCA data base and the flux densities at 3 mm available for 10 of them are larger than 400 mJy. Detections with total intensity $> 5\sigma$ are available in the Planck Legacy Catalogue (Planck Collaboration 2013) at 100, 217 and 353 GHz channels (roughly corresponding to ALMA band 3, 6 and 7 frequencies) in the position of 22 sources with median values above 450 mJy for all the bands. The median fractional polarization of the sample is above 4.8 per cent. The most polarized object is AT20G J210933–411020 with 1.9 Jy of total intensity and 10 per cent fractional polarization, flat spectrum, increasing fractional polarization with frequency, and only 14 per cent relative variability over 3 yr time. In the *Planck* channels the spectra become steep down to 428 mJy at 100 GHz and only 128 mJy at 217 GHz.

AT20G J063546–751616 has 5.33 Jy of total intensity flux density at 20 GHz, 6.2 per cent polarized, and remains above 1 Jy up to 1 mm frequencies. It is classified as a flat spectrum radio quasar at redshift $z = 0.653$. Several notes indicate the presence of a jet structure, but the 6 km visibilities identify it as point-like and modestly variable in ATCA observations over few years. Thanks to its position it is always visible to Southern hemisphere telescopes like ATCA and ALMA and stands as the most suitable polarization calibrator at high frequencies and low declinations.

8 CONCLUSIONS

We have conducted sensitive polarization and total intensity observations on a 20 GHz flux-limited sample of 189 objects selected in

Table 10. Candidate millimetric polarization calibrators (see the text for details on selection criteria). Columns are as follows: (1) AT20G name; (2) 18 GHz total intensity flux density and its error in mJy; (3) 18 GHz fractional polarization; (4–5) spectral indices $\alpha_{4,8}^{8,6}$ and $\alpha_{8,6}^{8,6}$ where low-frequency AT20G data are available; (6) variations of fractional polarization at different frequencies $\Delta\pi = \pi_{8,6\text{GHz}} - \pi_{18\text{GHz}}$ where low-frequency AT20G data are available; (7–8) epoch of AT20G catalogue run (see Murphy et al. 2010; ‘1’: 2004, ‘2’: 2005, ‘3–4’: 2006, ‘5–6’: 2007, ‘7’: 2008) and relative variation of flux densities ($S_{\text{AT20G}} - S_{\text{ex2006}}/S_{\text{AT20G}}$); (9) ratio between the visibilities amplitude averaged over the long baselines (~ 4.5 km) and the short ones (~ 0.2 km) in the AT20G catalogue (Chhetri et al. 2013); (10) 3 mm flux density extracted from the ATCA calibrator data base; (11–14) flux densities listed (above 5σ) in the Planck Legacy Catalogue (Planck Collaboration 2013) at 30, 100, 217 and 353 GHz, in mJy; (15–16) redshift and optical identification (‘QSO-G’: quasar and galaxy as classified in NED; ‘AeB’: AGN with broad line emissions; ‘Ae’: AGN with emission features; Mahony et al. 2011).

Name ‘AT20G’	S20 (mJy)	m20 (per cent)	$\alpha_{4,8}^{8,6}$	$\alpha_{8,6}^{8,6}$	$\Delta\pi$ (per cent)	Run	Variab 18 GHz	6 km Vis ratio	S_ATCAcal 3mm (mJy)	S_Planck 30 GHz	S_Planck 100 GHz	S_Planck 217 GHz	S_Planck 353 GHz	z	Opt id
J001035–302748	595.8 ± 11.9	5.23	0.26	-0.17	-1.48	1	-24.37	0.96	-	-	-	-	-	1.19	QSO
J004959–573827	1859.0 ± 37.2	3.84	0.16	-0.38	-1.11	2	-0.70	0.97	621.3	1262.19	651.72	324.88	-	1.797	QSO
J040658–382627	1627.0 ± 32.6	6.13	0.07	-0.26	-1.24	1	28.15	0.94	786.5	1402.08	925.25	552.50	374.80	1.285	QSO
J045550–461558	5333.0 ± 106.7	6.54	0.60	-0.11	-2.47	1	22.05	-	1204.6	2596.76	1407.34	840.84	516.70	0.853	QSO
J051644–620706	645.7 ± 12.9	4.57	0.14	-0.11	-0.41	2	-22.81	0.98	-	1183.03	633.39	630.04	379.25	-	BLLac
J055009–573224	976.7 ± 19.5	5.27	0.36	-0.17	0.06	2	-2.90	0.99	-	1171.74	818.22	599.80	505.13	2.001	QSO
J062307–643620	513.1 ± 10.3	3.33	-0.14	-0.26	-0.26	2	-39.74	0.97	432.4	835.80	546.79	393.24	390.74	0.129	G
J063546–751616	3311.0 ± 66.3	6.21	-0.17	-0.37	-2.19	4	5.10	0.99	2618	4295.48	3406.59	1831.33	1179.58	0.653	QSO
J074719–331046	1105.0 ± 22.1	3.55	0.39	0.37	-1.12	1	13.94	-	1283.7	-	1035.80	884.28	-	-	-
J081108–492943	544.9 ± 10.9	3.47	0.03	-0.30	1.2	1	-12.31	-	-	-	-	-	-	-	-
J102343–664647	519.9 ± 10.4	3.84	-0.16	-0.45	-3.09	2	-6.37	0.96	-	-	248.99	-	-	-	-
J110352–535700	555.8 ± 11.1	6.29	0.44	-0.57	-4.51	2	3.02	0.94	-	684.57	871.58	545.33	-	-	-
J111207–570339	643.4 ± 12.9	8.1	-0.02	-0.67	-0.71	2	-15.17	-	-	-	400.88	300.47	-	-	-
J111827–463415	857.6 ± 17.2	3.67	-0.02	-0.24	2.15	1	-11.59	-	-	1034.80	769.83	404.30	-	0.713	QSO
J115253–834410	521.0 ± 24.0	6.50	0.46	0.16	0.49	3	0.00	1	-	-	321.68	-	-	-	-
J121806–460030	687.6 ± 13.8	8.50	-0.55	-1.27	-3.03	1	-1.22	0.89	-	-	-	-	-	0.529	QSO
J142432–491349	2635.0 ± 52.8	6.94	-0.36	-0.76	-2.73	5	-0.27	0.76	1226	1942.63	992.69	616.51	-	-	-
J142756–420618	2495.0 ± 49.9	3.96	0.04	-0.40	-1.09	5	-9.98	0.98	980.8	1972.67	1898.62	1332.41	868.12	1.522	QSO
J170144–562155	1123.0 ± 22.5	6.06	-0.16	-0.73	-2.25	5	-3.38	1.01	500.7	1074.82	615.90	283.17	-	-	-
J170918–352520	568.9 ± 11.4	8.84	-0.31	-0.69	-3.82	1	-6.87	0.53	-	-	-	-	-	-	-
J171310–341827	908.2 ± 18.2	4.85	0.32	-0.49	-0.28	1	-12.97	1.01	-	-	-	-	-	-	-
J173315–372232	1001.0 ± 20.0	3.45	0.49	0.00	-1.71	5	-12.19	1.02	1270.6	-	1426.47	955.66	-	-	-
J182607–365102	558.6 ± 11.2	3.60	0.34	-0.15	-0.17	4	4.40	1.01	-	-	-	-	-	-	-
J194025–690757	596.7 ± 12.0	3.13	-0.13	-0.43	-2.41	2	12.52	0.98	-	-	-	-	-	3.154	QSO
J210933–411020	1905.0 ± 38.1	10.01	0.06	-0.49	-2.42	1	14.38	-	790.2	698.88	428.23	182.42	-	1.058	QSO
J214629–775554	1264.0 ± 2.0	3.02	0.91	0.23	-2.72	4	0.00	0.99	-	-	399.99	238.80	-	0.334	QSO
J220743–534633	1133.0 ± 22.7	4.25	-0.06	-0.40	1.41	2	0.97	0.95	-	-	548.39	305.55	-	1.206	QSO
J224616–560746	749.7 ± 15.0	8.07	0.43	-0.01	-2.89	2	-40.32	1	-	-	271.68	154.80	-	1.325	QSO
J232917–473019	1153.0 ± 23.1	3.97	-0.23	-0.41	-0.48	1	-23.07	0.96	1023.3	1578.01	1050.37	626.79	436.81	1.304	AeB

the AT20G survey, choosing sources that have $S_{20\text{GHz}} > 500$ mJy in the declination range $\delta < -30^\circ$ in the survey scans before 2006 October. They have been followed up during an observing run in 2006 October designed to reach 1 mJy sensitivity in polarization. This strongly improved the sensitivity and the detection rate for polarization observations over any previous sample investigated in this sky area at frequencies above 10 GHz.

94 per cent of the 180 extragalactic point sources have a detection of polarized flux density at least at 18 GHz. 172 of them have been observed also at 4.8 and 8.6 GHz, and 143 sources have a detected polarized flux density at all three frequencies.

The nine sources identified as extended have poor quality flux density measurement. So, for the sake of completeness, we extracted the values of polarized flux density from the 9 yr co-added *WMAP* maps. We recover an upper limit for five of them and a detection at 23 GHz for two of them (ForA and CenA). The final sample of 187 sources that we analysed constitutes a 99 per cent complete sample at the 2006 survey selection epoch with a 91.4 per cent polarization detection rate. In addition, detections have been obtained at all the *WMAP* frequencies for PicA.

This sample constitutes an ancillary data set for present and future studies of polarization in the Southern hemisphere and complements other samples recently observed either in equatorial regions (Sajina et al. 2011) or in the Northern hemisphere (Jackson et al. 2010). Analysis of the *WMAP* and *Planck* data (López-Caniego et al. 2009, and references therein) has demonstrated that similar source lists are crucial to improve the investigation of the CMB E and B modes in millimetric wavelength bands.

Thanks to our high detection rate, to a low polarized flux density level, to the multifrequency observations and to the inclusion of integrated flux densities for extended objects observed in mosaic mode with the ATCA (Burke-Spolaor et al. 2009) or extracted from the *WMAP* 9 yr maps (updating the findings of López-Caniego et al. 2009) the analysis of our sample in total intensity and polarization allowed us to draw the following conclusions.

(i) The spectral behaviours in total intensity and in polarization are different for any population of sources. This implies that it is extremely difficult to make an estimation of polarized flux densities from total intensity measurements.

(ii) There is no statistically significant evidence of increasing fractional polarization with frequency. This implies that Faraday depolarization is not strong enough to modify the spectral behaviour at and above ~ 4.8 GHz. Spectral behaviour in polarization is, in fact, the result of the combined effects of beam depolarization due to multiple components, chaotic magnetic fields and Faraday depolarization.

(iii) Thanks to our high detection rate we can state that there is no evidence of an anticorrelation of fractional polarization with total intensity flux density as was previously noted by several surveys, which were probably biased by a selection effect: only highly polarized sources can be detected for faint sources, while low fractional polarization percentages can be detected in bright objects; furthermore, faint objects in complete samples are typically more numerous than bright ones.

(iv) Thanks to the high sensitivity of our observations we were able to extend the polarization source counts at 18 GHz of Tucci & Toffolatti (2012) and to confirm their findings.

(v) We identified a list of 29 candidate calibrators for polarization at declination below -30° and frequencies $\gtrsim 20$ GHz. The best candidate is AT20G J063546–751616 that is point-like in the AT20G catalogue, has 5.33 Jy of total intensity flux density at 20 GHz,

is 6.2 per cent polarized and shows flux density above 1 Jy up to ~ 1 mm wavelengths.

ACKNOWLEDGEMENTS

MM, MN and GDZ acknowledge financial support for this research by ASI/INAF agreement I/072/09/0 for the *Planck* LFI activity of Phase E2. ML-C acknowledges partial financial support from the Spanish Ministerio de Economía y Competitividad project AYA-2012-39475-C02-01 and the Consolider Ingenio-2010 Programme project CSD2010-00064.

We thank the staff at the Australia Telescope Compact Array site, Narrabri (NSW), for the valuable support they have provided in running the telescope. The Australia Telescope Compact Array is part of the Australia Telescope which is funded by the Commonwealth of Australia for operation as a National Facility managed by CSIRO.

We thank Elisabetta Liuzzo (INAF-IRA) and Benjamin Walter (Haverford) for the useful discussions and proofreading. We thank the anonymous referee for the useful comments.

REFERENCES

- Agudo I., Thum C., Wiesemeyer H., Krichbaum T. P., 2010, *ApJS*, 189, 1
 Argüeso F., González-Nuevo J., Toffolatti L., 2003, *ApJ*, 598, 86
 Argüeso F., Sanz J. L., Herranz D., 2011, preprint (arXiv:1101.0701)
 Battye R. A., Browne I. W. A., Peel M. W., Jackson N. J., Dickinson C., 2011, *MNRAS*, 413, 132
 Blandford R. D., Königl A., 1979, *ApJ*, 232, 34
 Burke-Spolaor S., Ekers R. D., Massardi M., Murphy T., Partridge B., Ricci R., Sadler E. M., 2009, *MNRAS*, 395, 504
 Chhetri R., Ekers R. D., Mahony E. K., Jones P. A., Massardi M., Ricci R., Sadler E. M., 2012, *MNRAS*, 422, 2274
 Chhetri R., Ekers R. D., Jones P. A., Ricci R., 2013, *MNRAS*, 434, 956
 Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
 de Zotti G., Toffolatti L., Argüeso F., Davies R. D., Mazzotta P., Partridge R. B., Smoot G. F., Vittorio N., 1999, in Maiani L., Melchiorri F., Vittorio N., eds, *AIP Conf. Proc. Vol. 476, Conference on 3K Cosmology*. Am. Inst. Phys., New York, p. 204
 de Zotti G., Ricci R., Mesa D., Silva L., Mazzotta P., Toffolatti L., González-Nuevo J., 2005, *A&A*, 431, 893
 Gehrels N., 1986, *ApJ*, 303, 336
 Gold B. et al., 2011, *ApJS*, 192, 15
 Grant J. K., Taylor A. R., Stil J. M., Landecker T. L., Kothes R., Ransom R. R., Scott D., 2010, *ApJ*, 714, 1689
 Griffith M. R., Wright A. E., Burke B. F., Ekers R. D., 1994, *ApJS*, 90, 179
 Griffith M. R., Wright A. E., Burke B. F., Ekers R. D., 1995, *ApJS*, 97, 347
 Hancock P. J. et al., 2011, *Exp. Astron.*, 32, 147
 Jackson N., Battye R. A., Browne I. W. A., Joshi S., Muxlow T. W. B., Wilkinson P. N., 2007, *MNRAS*, 376, 371
 Jackson N., Browne I. W. A., Battye R. A., Gabuzda D., Taylor A. C., 2010, *MNRAS*, 401, 1388
 Klein U., Mack K.-H., Gregorini L., Vigotti M., 2003, *A&A*, 406, 579
 Kurinsky N., Sajina A., Partridge B., Myers S., Chen X., López-Caniego M., 2013, *A&A*, 549, A133
 López-Caniego M., Massardi M., González-Nuevo J., Lanz L., Herranz D., De Zotti G., Sanz J. L., Argüeso F., 2009, *ApJ*, 705, 868
 Mahony E. K. et al., 2011, *MNRAS*, 417, 2651
 Massardi M. et al., 2008, *MNRAS*, 384, 775
 Massardi M., López-Caniego M., González-Nuevo J., Herranz D., de Zotti G., Sanz J. L., 2009, *MNRAS*, 392, 733
 Massardi M. et al., 2011a, *MNRAS*, 412, 318
 Massardi M., Bonaldi A., Bonavera L., López-Caniego M., de Zotti G., Ekers R. D., 2011b, *MNRAS*, 415, 1597

- Mauch T., Murphy T., Buttery H. J., Curran J., Hunstead R. W., Piestrzynski B., Robertson J. G., Sadler E. M., 2003, *MNRAS*, 342, 1117
- Mesa D., Baccigalupi C., De Zotti G., Gregorini L., Mack K.-H., Vigotti M., Klein U., 2002, *A&A*, 396, 463
- Murphy T. et al., 2010, *MNRAS*, 402, 2403
- Planck Collaboration, 2011a, *A&A*, 536, A7
- Planck Collaboration, 2011b, *A&A*, 536, A13
- Planck Collaboration, 2013, preprint (arXiv:1303.5088)
- Ricci R., Prandoni I., Gruppioni C., Sault R. J., De Zotti G., 2004, *A&A*, 415, 549
- Sadler E. M. et al., 2006, *MNRAS*, 371, 898
- Sajina A., Partridge B., Evans T., Steff S., Vechik N., Myers S., Dicker S., Korngut P., 2011, *ApJ*, 732, 45
- Sault R. J., Teuben P. J., Wright M. C. H., 1995, in Shaw R. A., Payne H. E., Hayes J. J. E., eds, *ASP Conf. Ser. Vol. 77, Astronomical Data Analysis Software and Systems IV*. Astron. Soc. Pac., San Francisco, p. 433
- Sault R. J., Hamaker J. P., Bregman J. D., 1996, *A&AS*, 117, 149
- Sault R. J., Rayner D. P., Kesteven M. J., 2002, in *AIP Conf. Proc. Vol. 609, Astrophysical Polarized Backgrounds*. Am. Inst. Phys., New York, p. 150
- Simmons J. F. L., Stewart B. G., 1985, *A&A*, 142, 100
- Subrahmanyan R., Ekers R. D., Saripalli L., Sadler E. M., 2010, *MNRAS*, 402, 2792
- Taylor A. R. et al., 2007, *ApJ*, 666, 201
- Toffolatti L., Argüeso Gomez F., de Zotti G., Mazzei P., Franceschini A., Danese L., Burigana C., 1998, *MNRAS*, 297, 117
- Toffolatti L., De Zotti G., Argüeso F., Burigana C., 1999, in de Oliveira-Costa A., Tegmark M., eds, *ASP Conf. Ser. Vol. 181, Microwave Foregrounds*. Astron. Soc. Pac., San Francisco, p. 153
- Tucci M., Toffolatti L., 2012, *Adv. Astron.*, 2012, 624987
- Tucci M., Martinez-Gonzalez E., Toffolatti L., Gonzalez-Nuevo J., De Zotti G., 2004, *MNRAS*, 349, 1267
- Wardle J. F. C., Kronberg P. P., 1974, *ApJ*, 194, 249
- Wright E. L. et al., 2009, *ApJS*, 180, 283

SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

Table 2. Catalogue of the sources of the extragalactic sample observed in the AT20G run dedicated to high sensitivity polarization (<http://mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stt1717/-/DC1>).

Please note: Oxford University Press are not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.